

Testing gravity with large scale structure dynamics

Elise Jennings
University of Chicago

Outline

- Accelerating expansion
- Growth of structure & Expansion history
- Redshift space distortions :
an important cosmological probe
- Measuring the growth rate:
Dark energy vs. modified gravity

EJ, B. Li, C.M. Baugh, G. Zhao, K. Kazuya 2013

EJ 2012

EJ, C. M. Baugh, S. Pascoli 2011

Probing Dark Energy

Supernovae, CMB BAO,
Weak Lensing & Cluster
observations



Λ CDM cosmology

$$\Omega_{\text{TOTAL}} \sim 1$$

~25% Dark Matter

~75% Dark Energy

What is causing the accelerating expansion ?

Is dark energy constant or does it evolve with time?

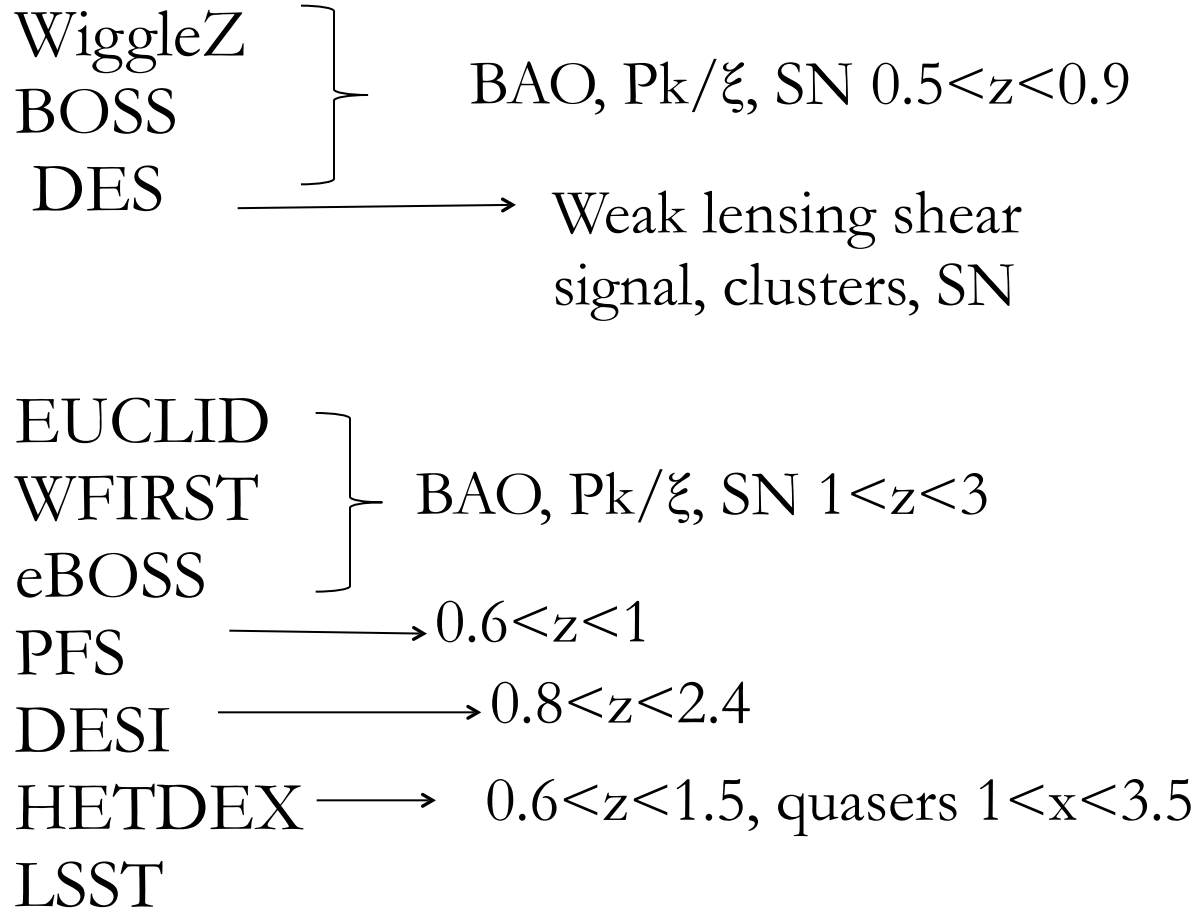
General Relativity or Modified gravity?

Missions to study the dark sector need precise measurements:

Expansion history $H(z)$

Growth rate $f = d \ln D / d \ln a$

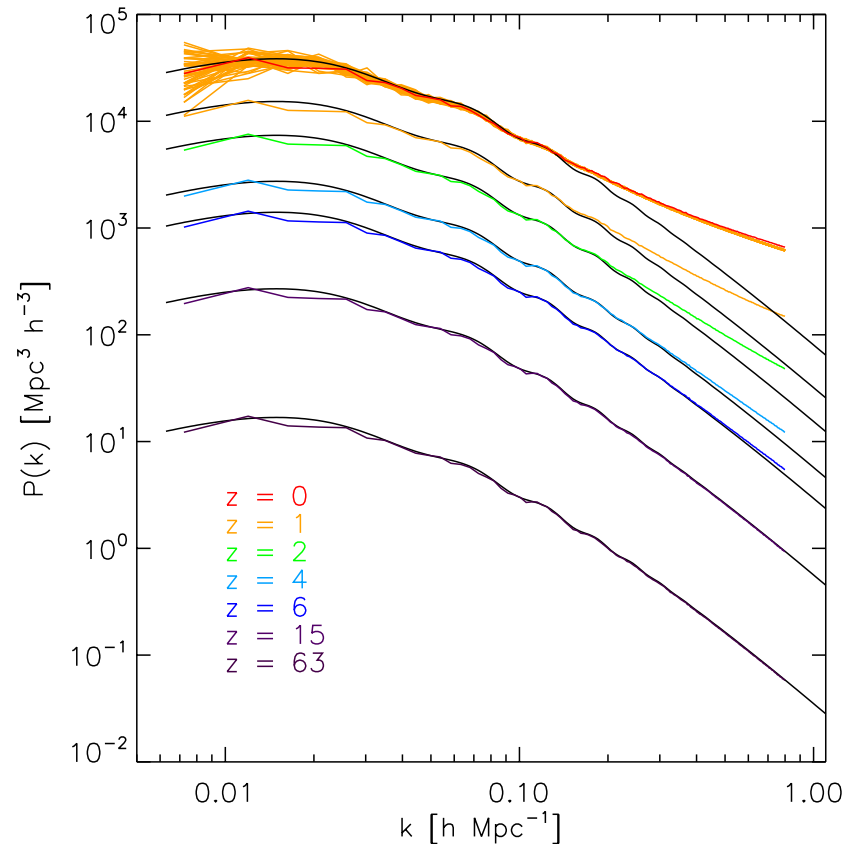
Probing Dark Energy



Goal:
Measure the
growth rate to
1% over
 $0.25 < z < 3$
Expansion rate
to <1%

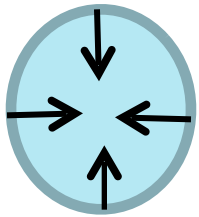
N-body simulations are an essential tool

- Model **nonlinear fluctuation growth, peculiar motions, nonlinear and scale dependent bias**
- Study nonlinear growth in different cosmologies

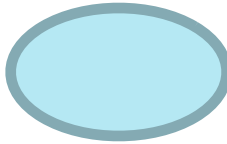


Redshift space distortions

Real space

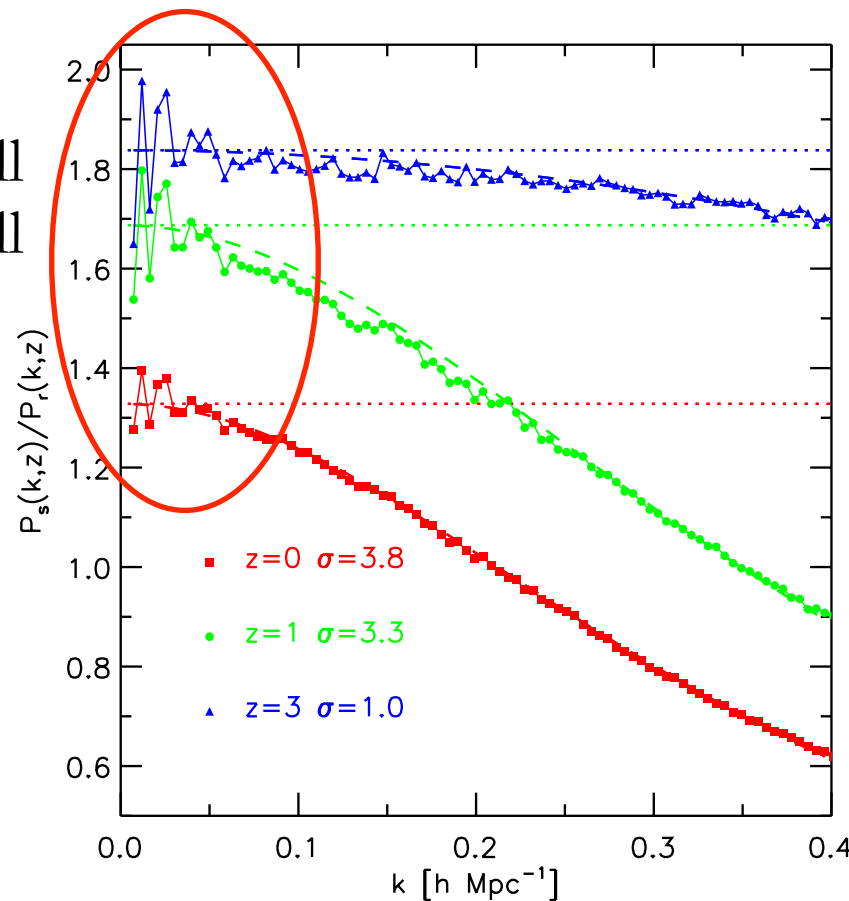


Redshift space



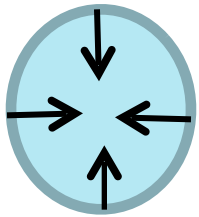
$$s = r + v$$

Large scales: Galaxies undergoing infall towards a spherical overdensity : shell appears squashed.

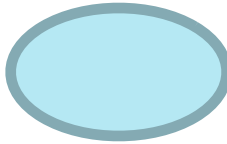


Redshift space distortions

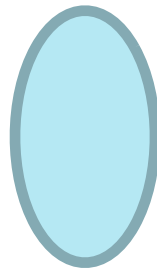
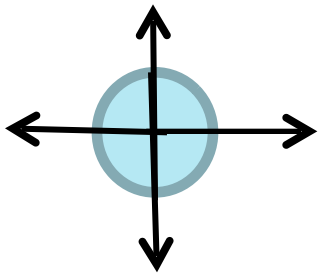
Real space



Redshift space

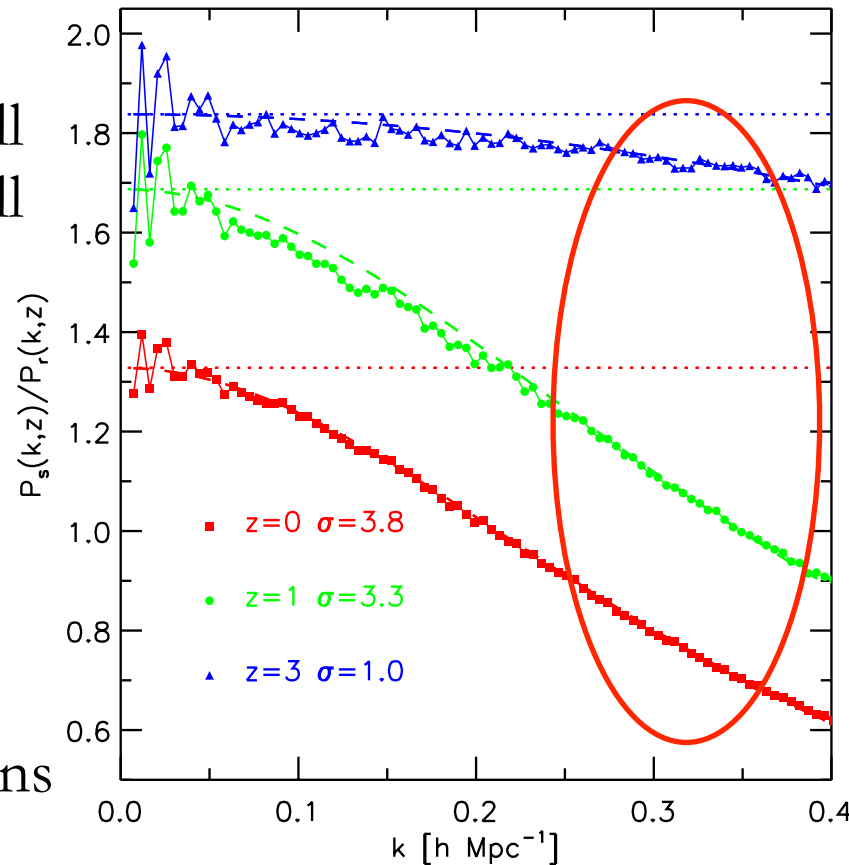


Large scales: Galaxies undergoing infall towards a spherical overdensity : shell appears squashed.



Small scales: Incoherent peculiar motions giving rise to “fingers-of-god”.

$$s = r + v$$



Clustering statistics

The Dark Energy Science report aka “Rocky III” :

RSD “among the most powerful ways of addressing whether the acceleration is caused by dark energy or modified gravity”

$$\delta_D(\vec{k}) + P(k, \mu) = \int \frac{d^3r}{(2\pi)^3} e^{-i\vec{k}\cdot\vec{r}} \langle e^{ik\mu\Delta u_z} [1 + \delta(\vec{x})][1 + \delta(\vec{x}')]\rangle$$

Power spectrum in redshift space

$v_z(\vec{x}) - v_z(\vec{x}')$

Linear theory

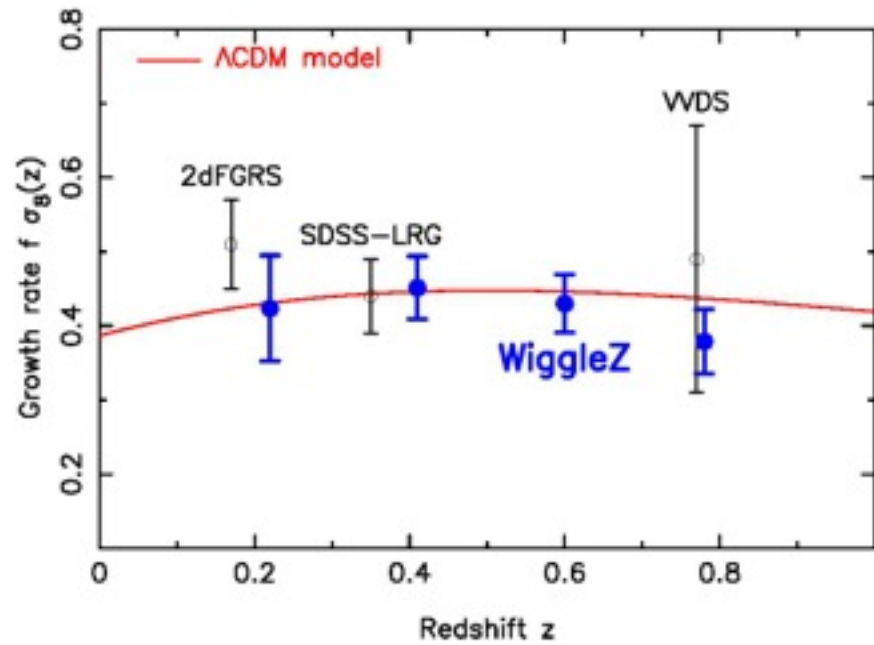
$$\delta_s(r) = \delta_r(r)(1 + \mu^2\beta)$$

Growth rate

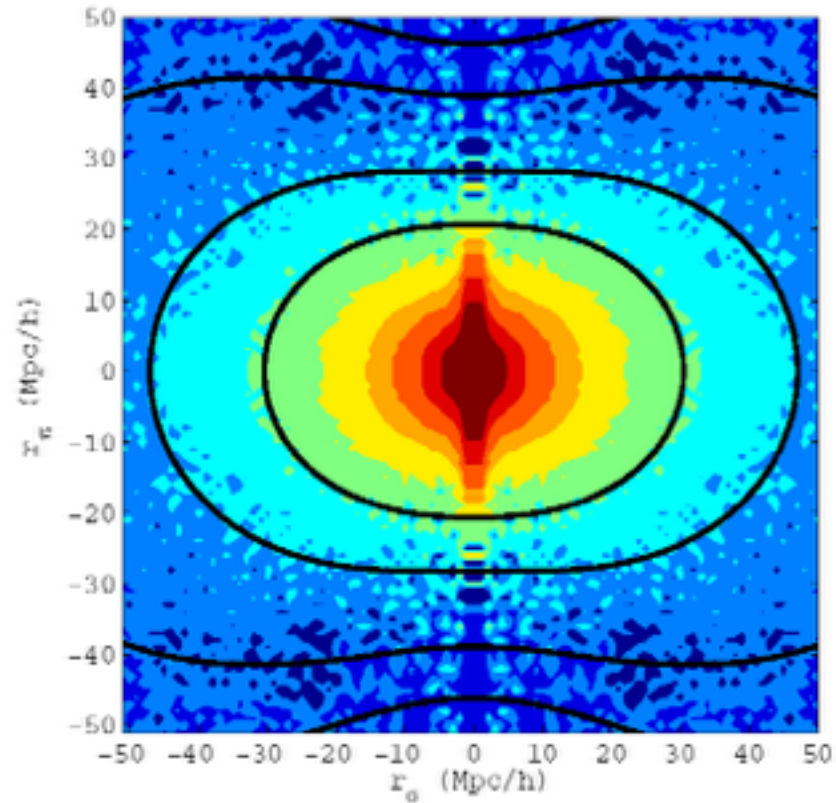
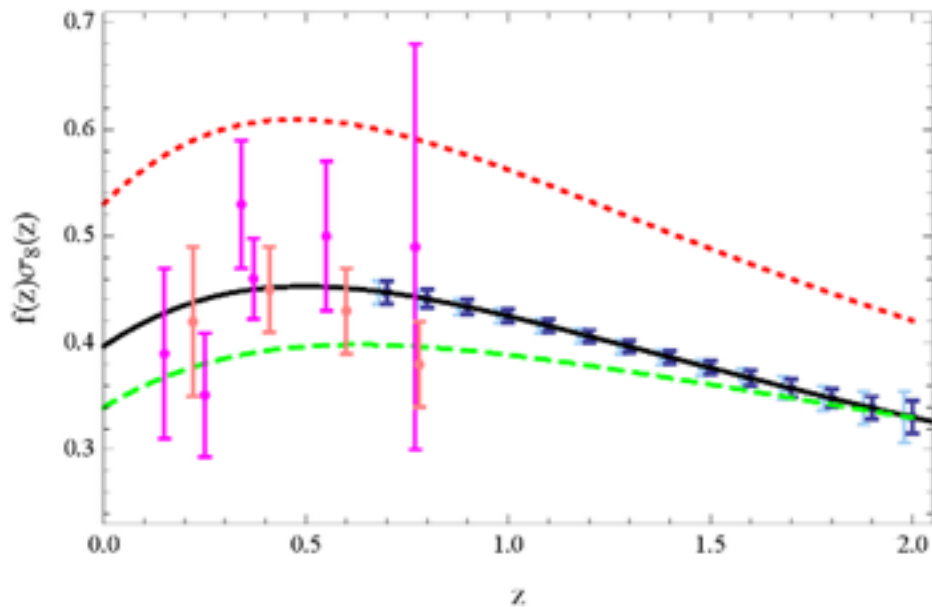
$$\beta = f/b$$

Linear bias

Current measurements & forecasts



Blake et al 2011



$$z=0.57: \quad f\sigma_8(a) = 0.43 \pm 0.07$$

Reid et al. 2012

Majerotto et al. 2012

Clustering statistics

Power spectrum in linear theory

$$P_s(k, \mu) = P_r(k) (1 + \mu^2 \beta)^2$$

$$P_l^s(k) = \frac{2l+1}{2} \int_{-1}^1 P(k, \mu) L_l(\mu) d\mu$$

$$\beta = f/b$$

Growth rate

Linear bias

$$\begin{pmatrix} P_0(k) \\ P_2(k) \end{pmatrix} = P_m(k) \begin{pmatrix} 1 + \frac{2}{3}\beta + \frac{1}{5}\beta^2 \\ \frac{4}{3}\beta + \frac{4}{7}\beta^2 \end{pmatrix} \quad \text{Kaiser (1987)}$$

Correlation function

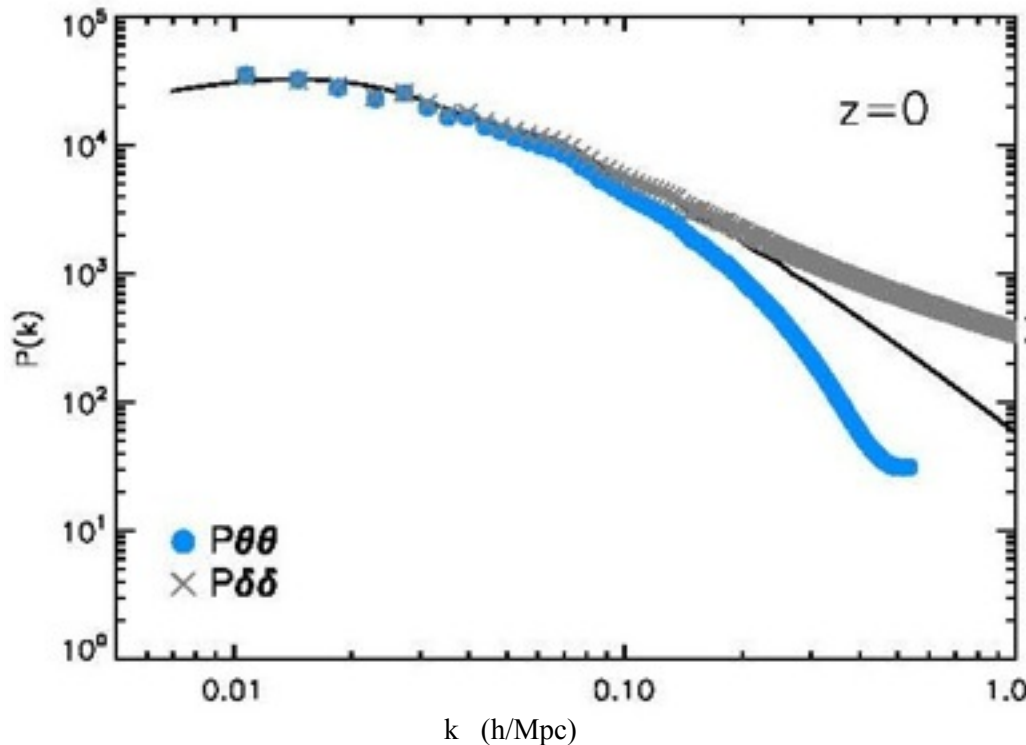
$$\begin{pmatrix} \xi_0(s) \\ \xi_2(s) \end{pmatrix} = \begin{pmatrix} (1 + \frac{2}{3}\beta + \frac{1}{5}\beta^2) \xi(r) \\ (\frac{4}{3}\beta + \frac{4}{7}\beta^2) [\xi(r) - \bar{\xi}(r)] \end{pmatrix}$$

Improved model for the redshift space $P(k)$:

$$P^s(k, \mu) = (P_{\delta\delta} + 2\mu^2 P_{\delta\theta} + \mu^4 P_{\theta\theta})$$

Velocity divergence cross $P(k)$

Velocity divergence auto $P(k)$



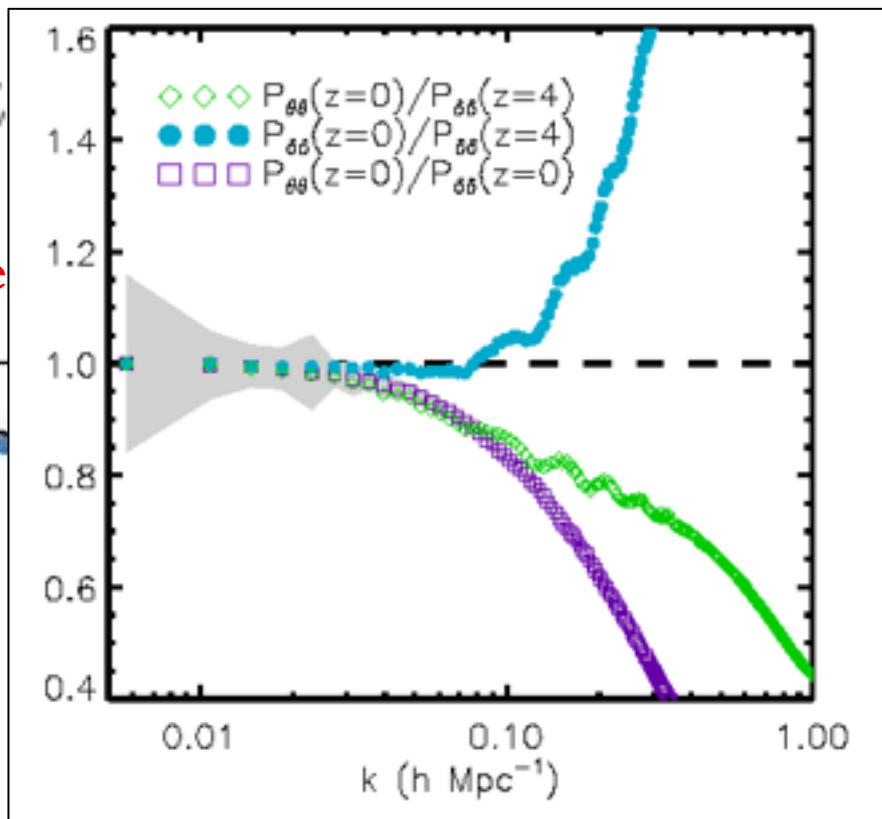
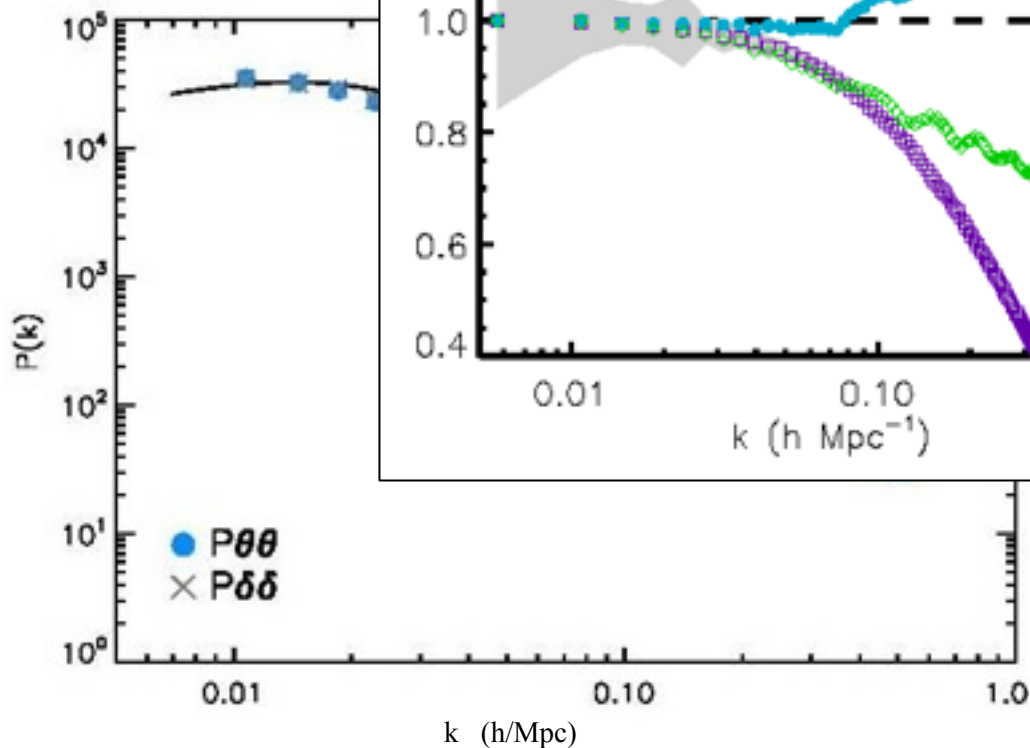
On what scales is the linear continuity equation valid ?

$$\vec{\nabla} \cdot \vec{v} = \theta = -H f(a) \delta(a)$$

Improved model for the redshift space $P(k)$:

$$P^s(k, \mu)$$

Velocity divergence



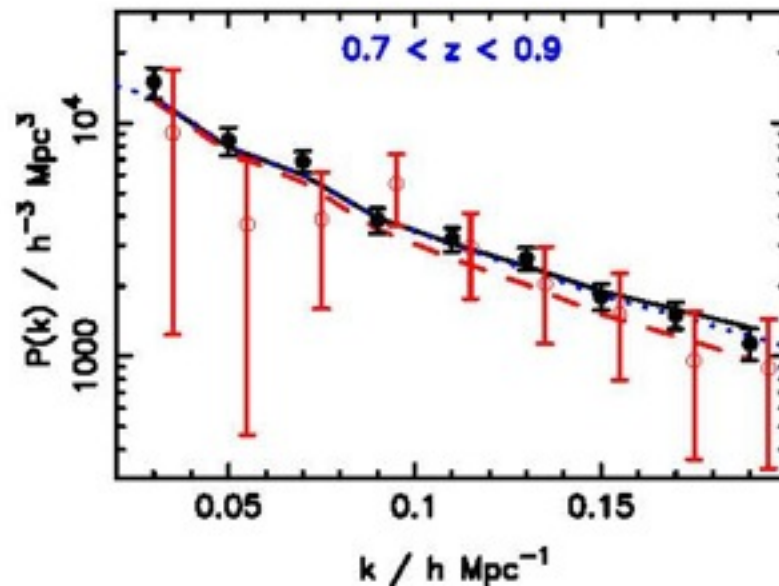
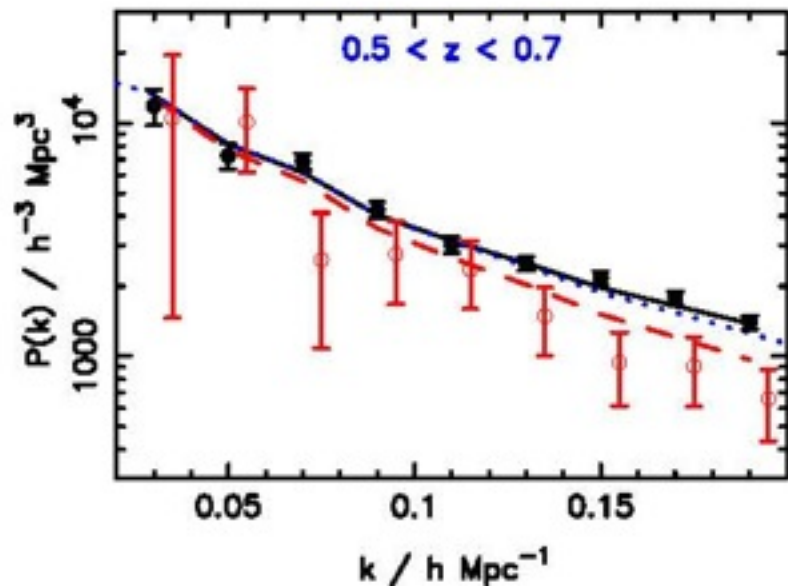
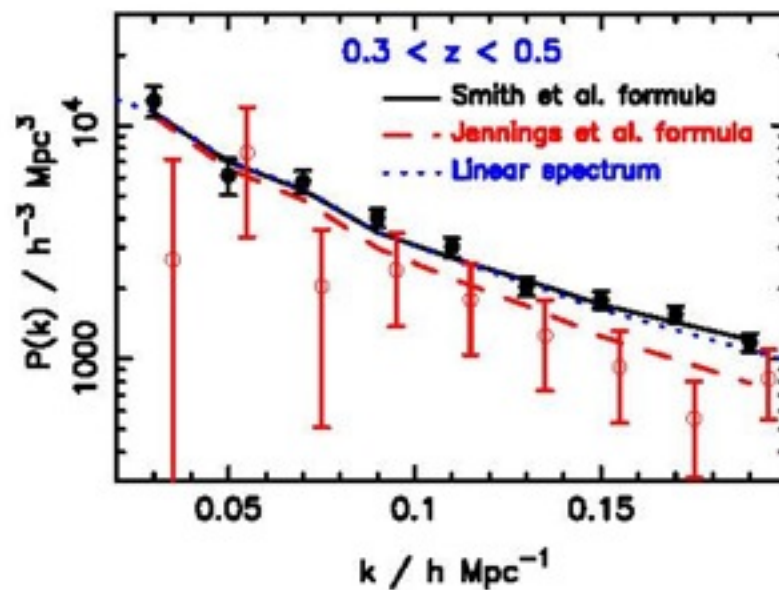
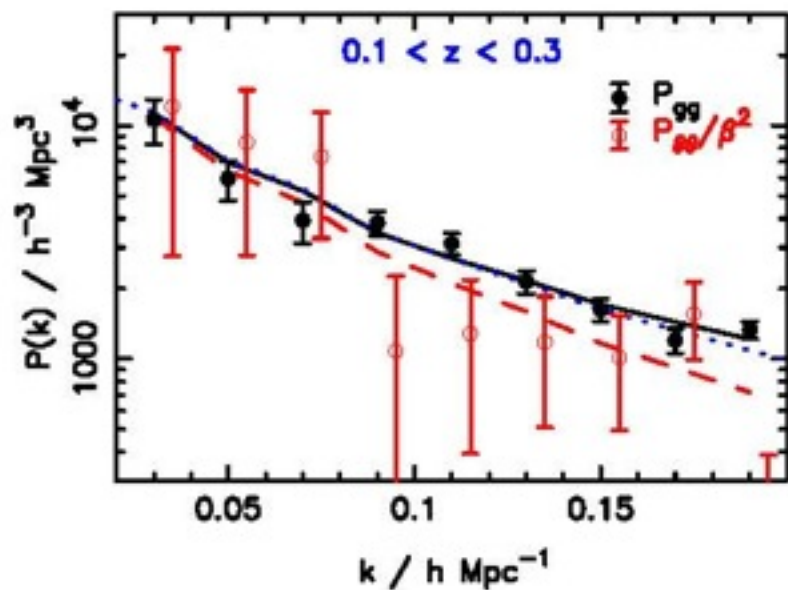
$$\mu^4 [P_{\theta\theta}]$$

convergence auto $P(k)$

at scales is the linear continuity equation valid?

$$\theta = -H f(a) \delta(a)$$

Results from WiggleZ



Dark Energy or modified gravity

Anything that can simultaneously explain

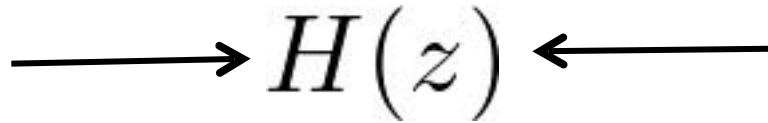
- **Angular diameter distances** (BAO, CMB)
- **Luminosity distances** (Supernovae Ia)

can be called “dark energy” but could be the result of modified gravity!

Measuring the **expansion history alone will not distinguish modified gravity from smooth dark energy**

Need to break the degeneracy with measurements of **growth factor**

Dark energy



$$H(z)$$

Modified gravity

e.g Quintessence

e.g. parametrised by
 $\mu^2 = G/G_N$ and
 $\zeta = 1 - \Psi/\Phi$ “slip parameter”

Dark energy $\longrightarrow H(z)$ \longleftarrow Modified gravity

e.g Quintessence



Within GR growth of density perturbations grows according to

$$\ddot{\delta} + 2H\dot{\delta} - 4\pi G_N \rho_m \delta = 0$$

e.g. parametrised by
 $\mu^2 = G/G_N$ and
 $\zeta = 1 - \Psi/\Phi$ “slip parameter”



If gravitational constant varies & $g = \delta/a$

$$\delta_D(\vec{k}) + P(k, \mu) = \int \frac{d^3r}{(2\pi)^3} e^{-i\vec{k}\cdot\vec{r}} \langle e^{ik\mu\Delta u_z} [1 + \delta(\vec{x})][1 + \delta(\vec{x}')] \rangle$$

Clustering statistics

Dark energy

$$\longrightarrow H(z)$$

Modified gravity

e.g Quintessence



e.g. parametrised by $\mu^2 = G/G_N$ and $\zeta = 1 - \Psi/\Phi$ “slip parameter”



Within GR growth of density perturbations grows according to

$$\ddot{\delta} + 2H\dot{\delta} - 4\pi G_N \rho_m \delta = 0$$

If gravitational constant varies & $g = \delta/a$

$$\delta_D(\vec{k}) + P(k, \mu) = \int \frac{d^3r}{(2\pi)^3} e^{-i\vec{k}\cdot\vec{r}} \langle e^{ik\mu\Delta u_z} [1 + \delta(\vec{x})][1 + \delta(\vec{x}')] \rangle$$

Consistency test of cosmology : compare growth & expansion rates

Growth rate

$$f(a) = \frac{d \ln \delta}{d \ln a}$$

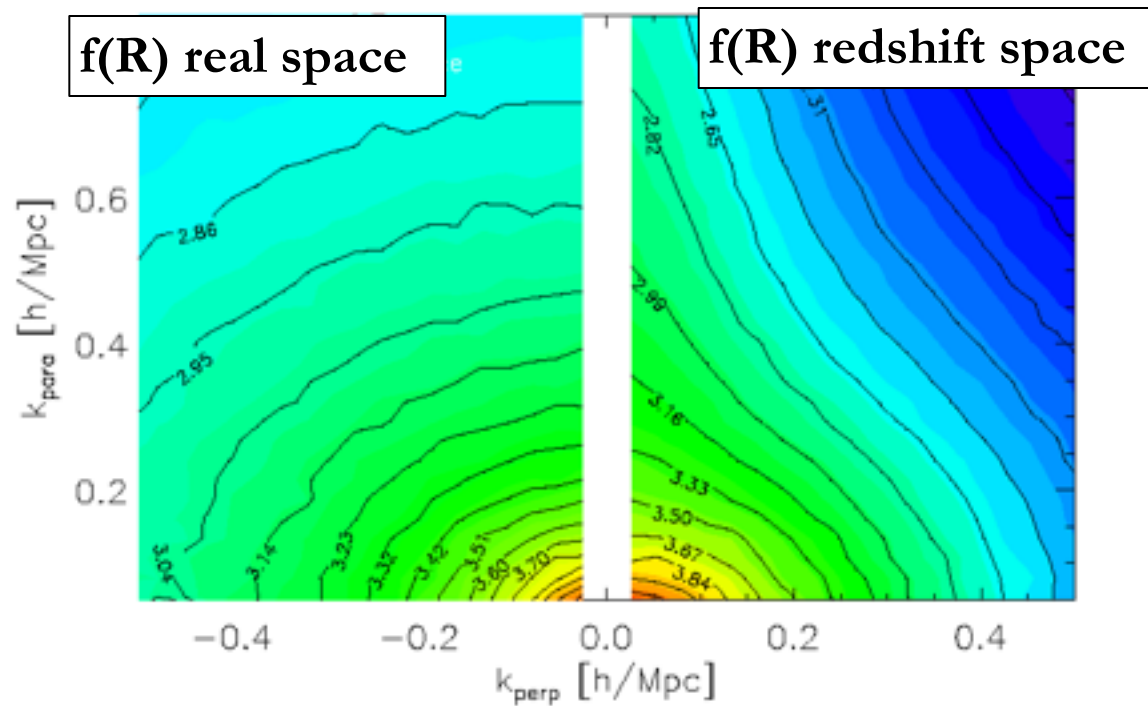
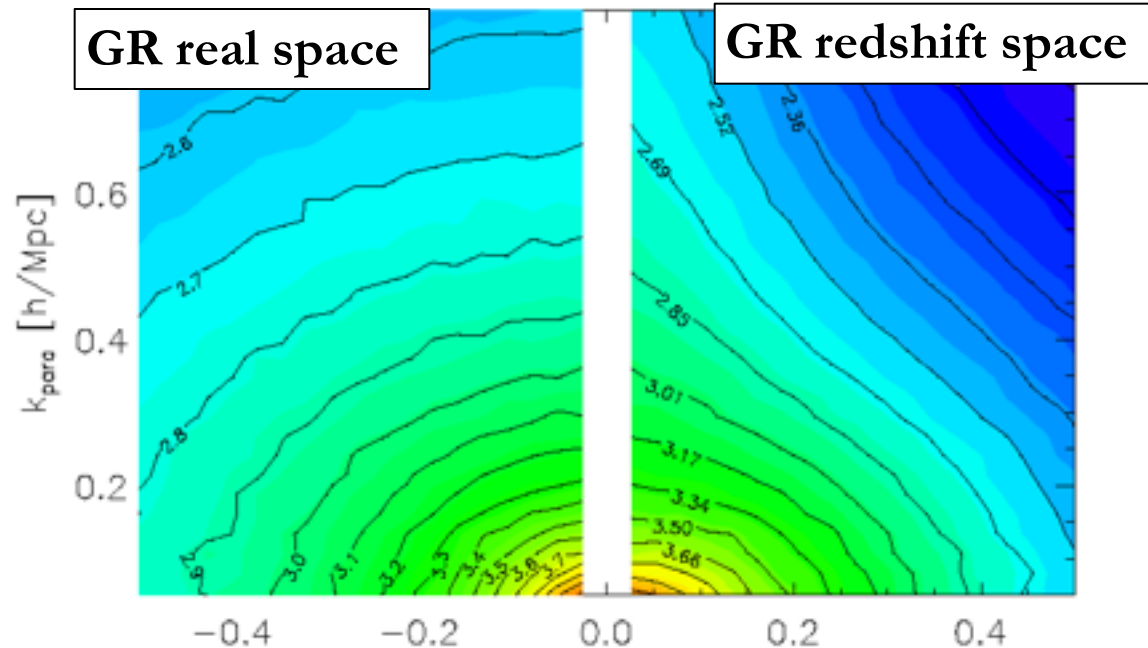
Clustering statistics

Growth rate

$$f(a) = \frac{d \ln \delta}{d \ln a}$$

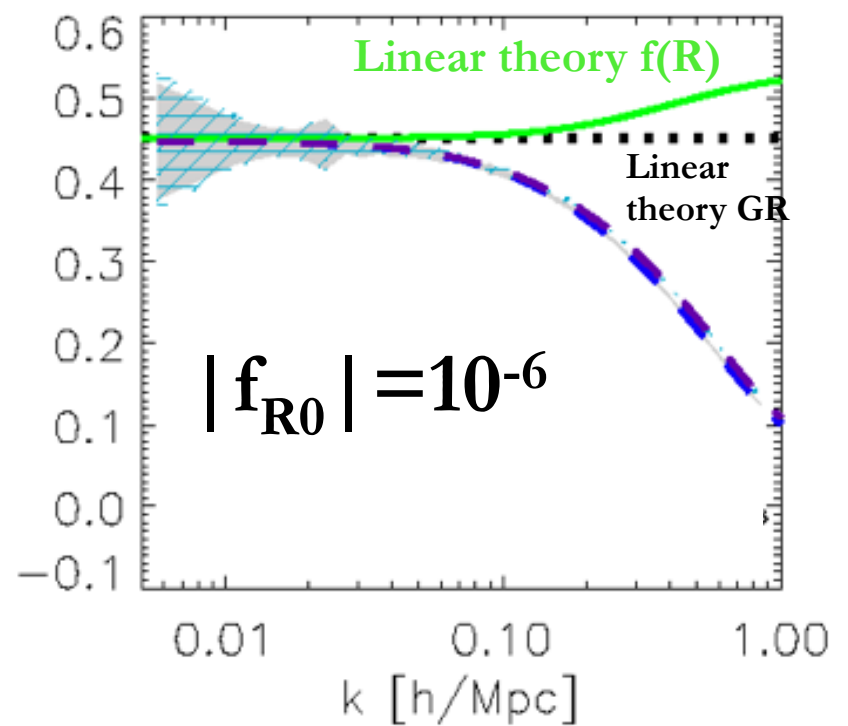
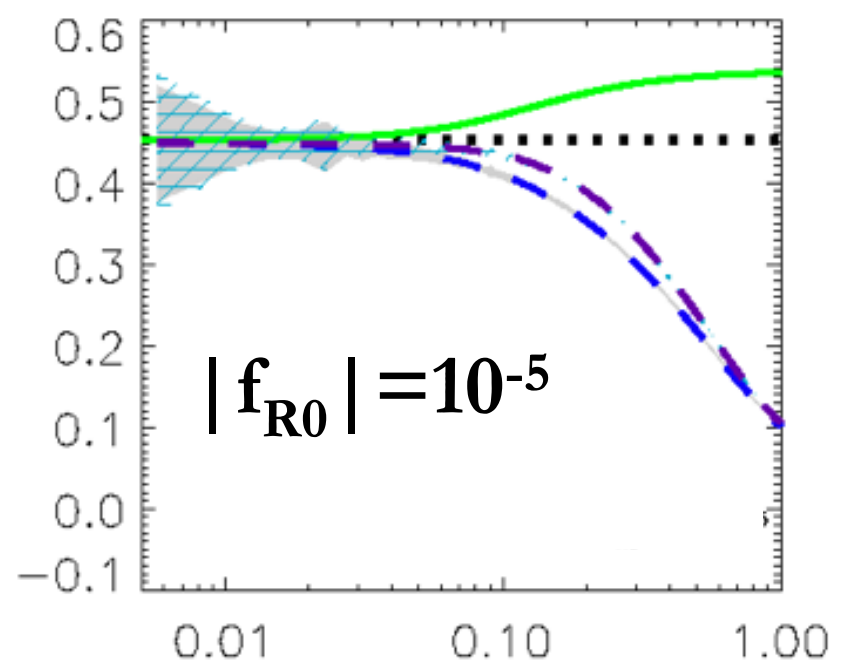
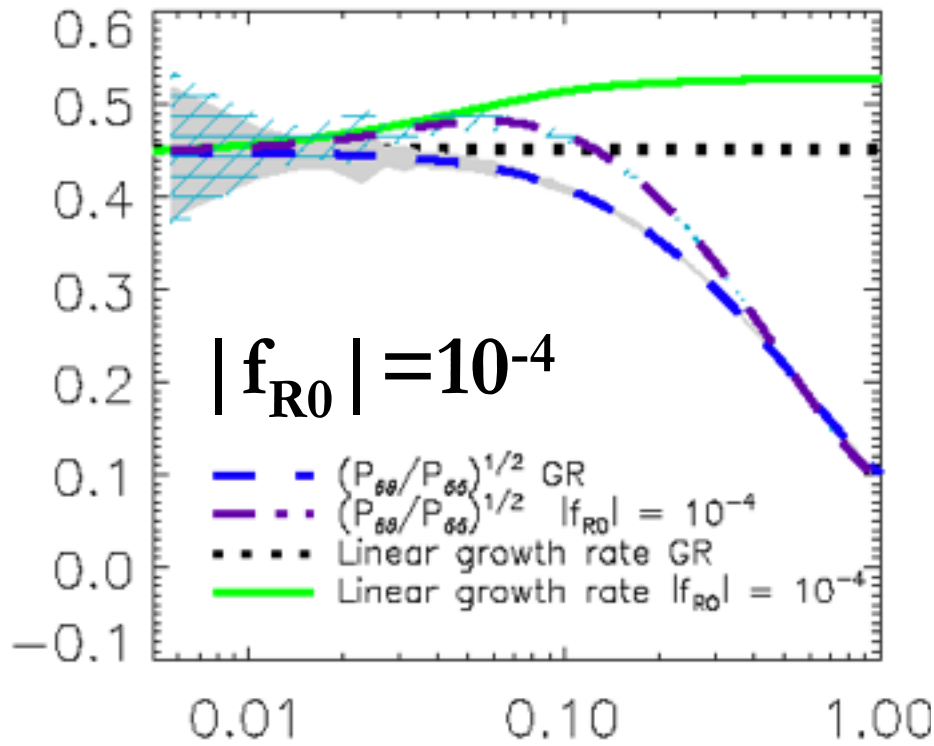
RSD in $f(R)$ gravity

Simulations:
Modified Ramses



EJ, Baugh, Li,
Zhao & Koyama
2012

\underline{P}_{00}

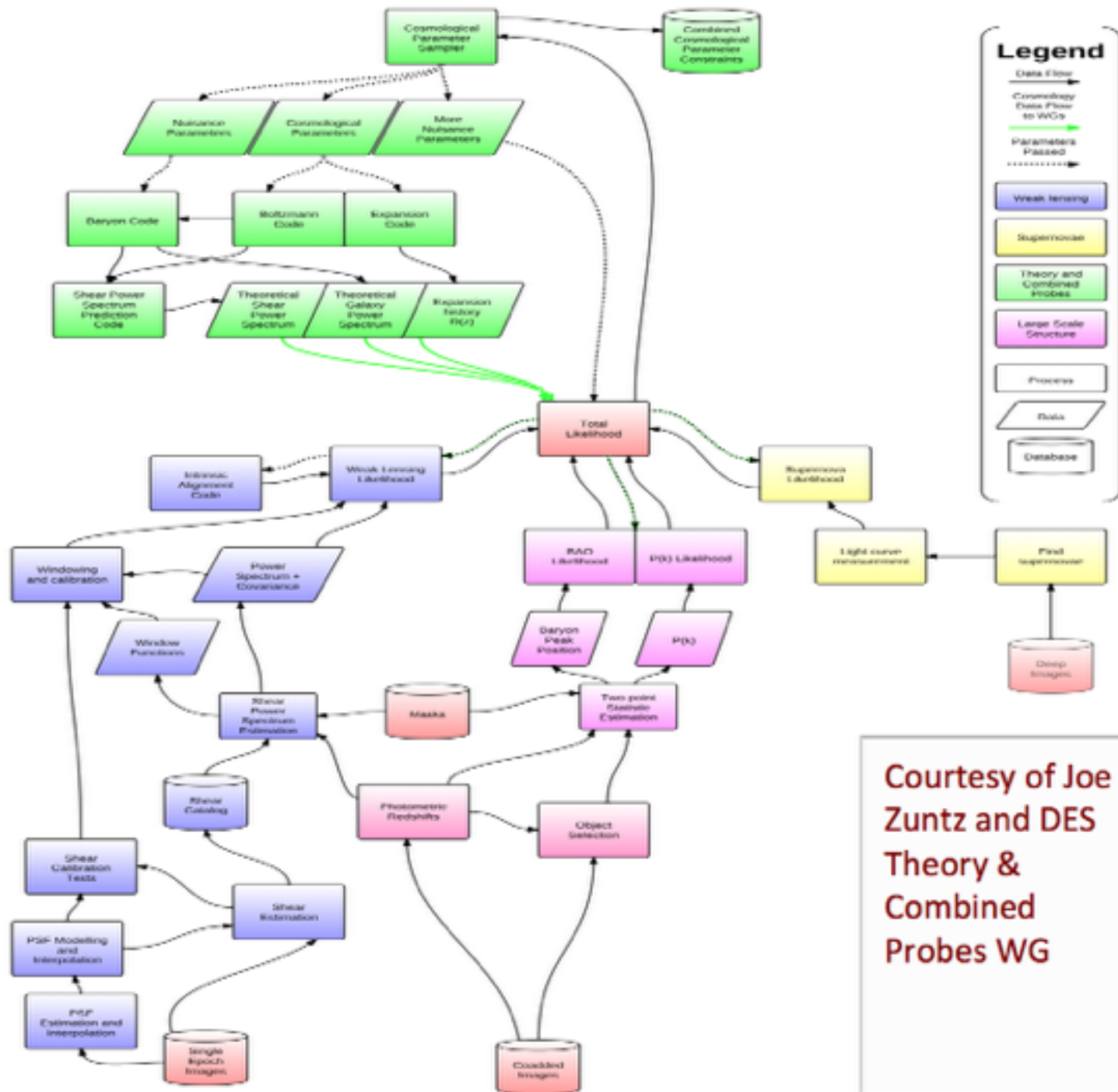


$$\vec{\nabla} \cdot \vec{v} = \theta = -H f(a) \delta(a)$$

CosmoSIS: modular cosmological parameter estimation

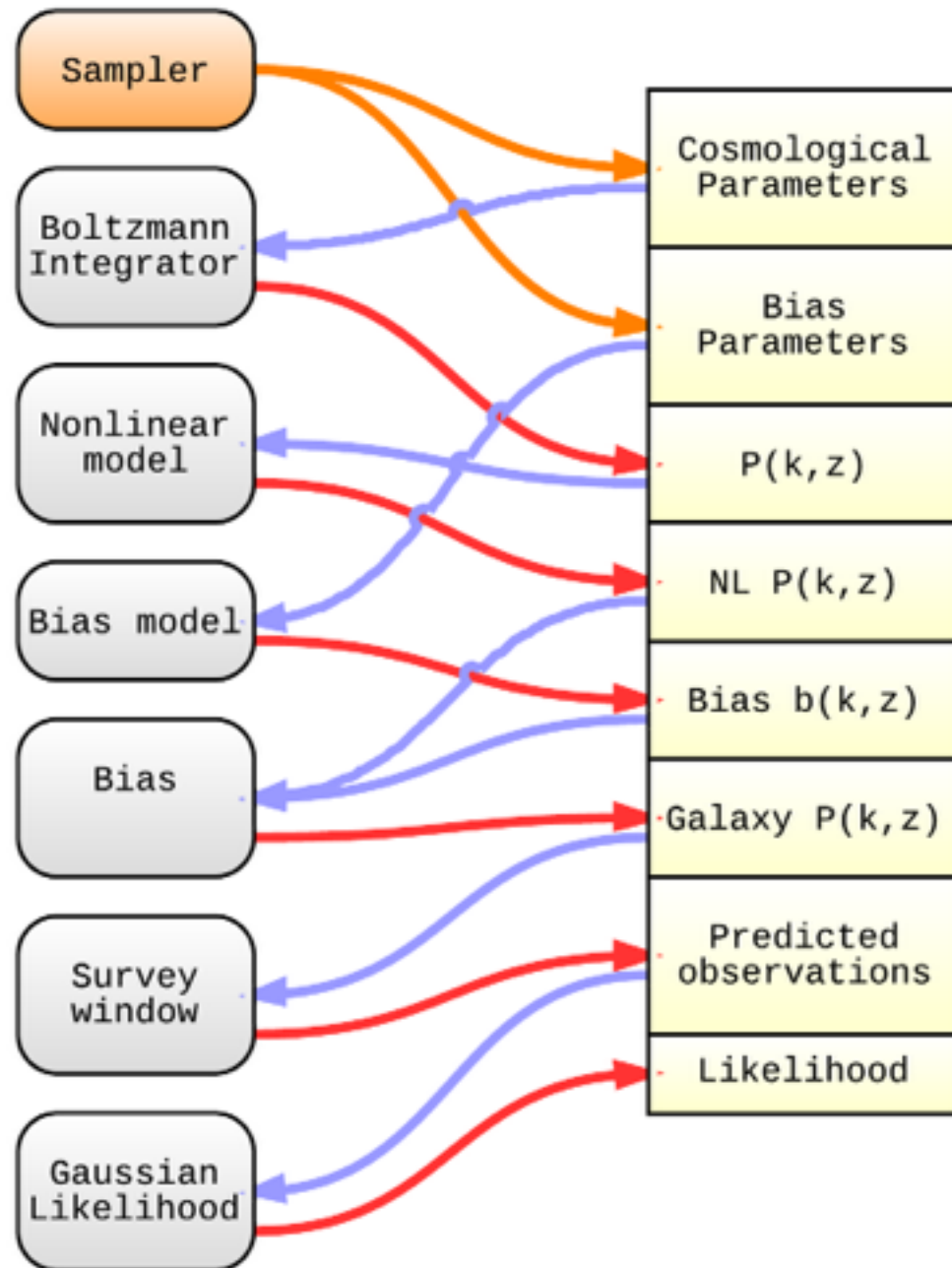
**Sarah Bridle, Scott Dodelson, Elise Jennings, Jim Kowalkowski,
Alessandro Manzotti, Marc Paterno, Doug Rudd, Saba Sehrish, Joe
Zuntz**





Modular Calculations

Zuntz et al 2014



CosmoSIS

CosmoSIS Standard Library
CAMB, Planck, WMAP,
BICEP2, BOSS, CRL...

DES specific
modules

LSST specific
modules

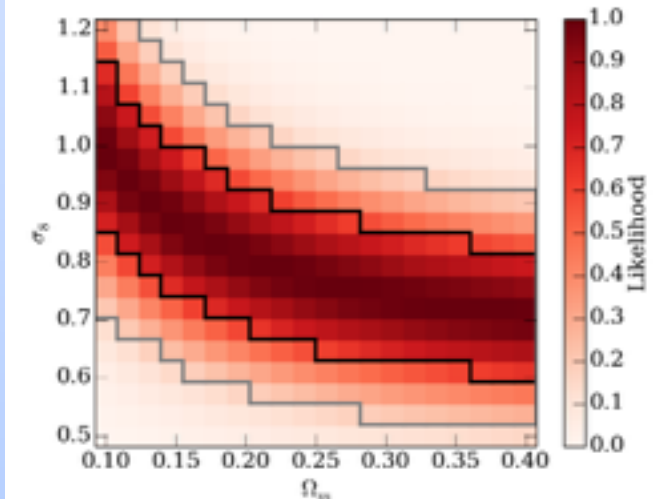
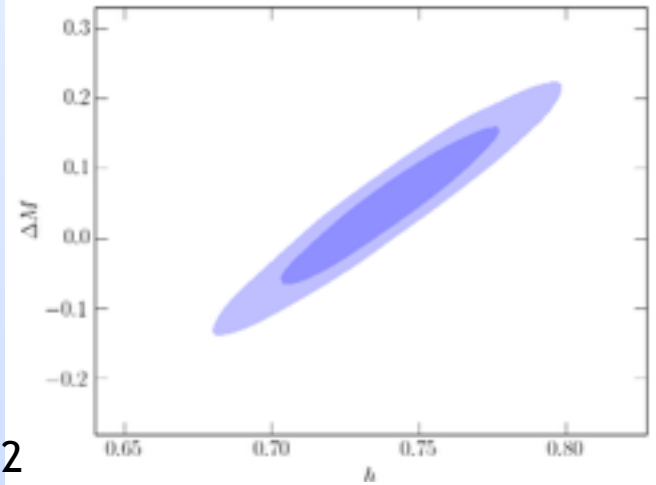
Other user's
physics
modules

CosmoSIS:
-core libraries
-infrastructure for
modules to interact

Software tools:
gcc, g++,
gfortran, Python,
SciPy, fftw,
gsl, NumPy, cfitsio,
pyfits

Key highlights

- easy download and install of all packages
- access to 3rd party codes - CAMB, Planck, WMAP, CRL BICEP2
- End User can - contribute to CosmoSIS Standard Library
 - use other's code
 - get credit for their work
 - add modules in C, C++, fortran, Python
- Provenance tracking



<https://bitbucket.org/joezuntz/cosmosis/wiki/Home>